

Spherical collapse in Modified Gravity theories

Paul de Fromont

supervised by *Jean-Michel Alimi*

July 5, 2016

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- accelerated cosmic expansion ! $\ddot{a} > 0$

- Dark Energy

- cosmological constant Λ

- $p = -\rho$

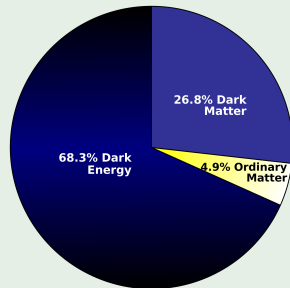
- dark fluid

- $p = w\rho$ and $w < -1/3$

- quintessence

- $p = w(a)\rho$ with $w(1) \sim -1$

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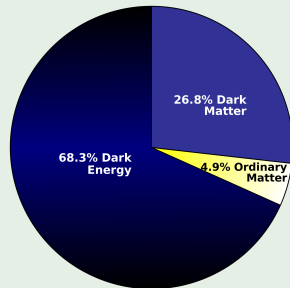
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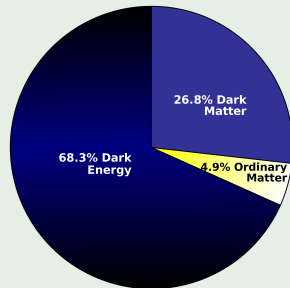
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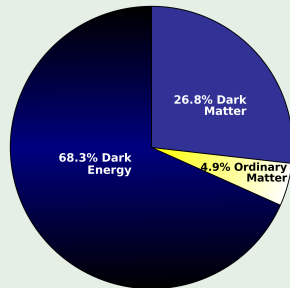
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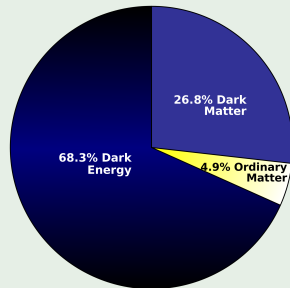
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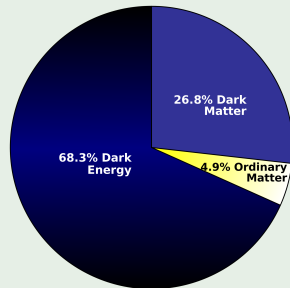
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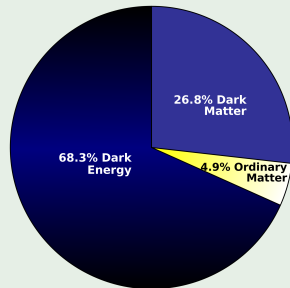
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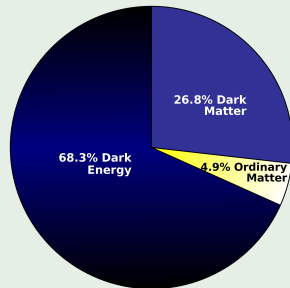
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dark energy : new exotic fluid

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = \kappa T_{\mu\nu} \quad (1)$$

modified gravity : modifications of the geometrical term

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = \kappa T_{\mu\nu} \quad (2)$$

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From the action point of view, Usual GR, introducing a new scalar field ϕ , minimally coupled to $g_{\mu\nu}$ (Quintessence)

$$S = \frac{1}{2\kappa} \int d^4x \sqrt{-g} R + S_m[\Psi; g] + S_{DE}[\phi, g] \quad (3)$$

leads to w :

$$w(\phi) = \frac{\dot{\phi}^2 - 2V(\phi)}{\dot{\phi}^2 + 2V(\phi)} \quad (4)$$

if $\dot{\phi}^2 \ll V(\phi)$, then $w \sim -1$!

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Non minimal coupling ? changing

$$S = \frac{1}{2\kappa} \int d^4x \sqrt{-g} R + S_m[\Psi; g] + S_{DE}[\phi, g] \quad (5)$$

to a non minimal coupling

$$S = \frac{1}{2\kappa} \int d^4x \sqrt{-g} F(\phi) R + S_m[\Psi; g] + S_{DE}[\phi, g] \quad (6)$$

results into

$$G_{\mu\nu} = \kappa [T_{\mu\nu}^m + T_{\mu\nu}(\phi)] \rightarrow F(\phi) G_{\mu\nu} = \kappa [T_{\mu\nu}^m + T_{\mu\nu}(\phi)] \quad (7)$$

known as Scalar-Tensor theories (**Modified Gravity**)

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Modified Gravity

The background dynamic

What is the background dynamics ? Starting from the Einstein equation

$$F(\phi)G_{\mu\nu} = \kappa T_{\mu\nu}^m + \kappa T_{\mu\nu}(\phi) \quad (8)$$

FLRW metric $ds^2 = dt^2 - a^2(t)d\mathbf{x}^2$,

$$F(a)G_{\mu\nu}(a) = \kappa T_{\mu\nu}^m(a) + \kappa T_{\mu\nu}(a) \quad (9)$$

can be directly transposed to

$$G_{\mu\nu}(a) = \kappa T_{\mu\nu}^m(a) + \underbrace{\kappa T_{\mu\nu}(a) + G_{\mu\nu}(a)(1 - F(a))}_{\kappa \tilde{T}_{\mu\nu}(a)} \quad (10)$$

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What is the background dynamics ? Starting from the Einstein equation

$$F(\phi)G_{\mu\nu} = \kappa T_{\mu\nu}^m + \kappa T_{\mu\nu}(\phi) \quad (8)$$

FLRW metric $ds^2 = dt^2 - a^2(t)d\mathbf{x}^2$,

$$F(a)G_{\mu\nu}(a) = \kappa T_{\mu\nu}^m(a) + \kappa T_{\mu\nu}(a) \quad (9)$$

can be directly transposed to

$$G_{\mu\nu}(a) = \kappa T_{\mu\nu}^m(a) + \underbrace{\kappa T_{\mu\nu}(a) + G_{\mu\nu}(a)(1 - F(a))}_{\kappa \tilde{T}_{\mu\nu}(a)} \quad (10)$$

For the **background** : *Modified Gravity = Dark Energy*

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What happens for the perturbed metric ? Consider

$$ds^2 = dt^2(1 + 2\Phi) - a^2 dx^2(1 - 2\Psi) \quad (11)$$

$$\rho_m = \bar{\rho}_m(1 + \delta) \quad (12)$$

$$\phi = \bar{\phi} + \delta\phi \quad (13)$$

and using

- Einstein equations
- scalar field *e.o.m.*
- assume that $\partial_t^2 \ll \nabla^2$ and $k \gg H = \dot{a}/a$

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2 in GR, $\mu(a, k) = G$

3 modified slip equation : $\Psi = \gamma(a, k)\Phi$

4 in GR = $\gamma = 1$

with

$$\mu = \frac{1 + p_3 k^2}{p_4 + p_5 k^2}, \quad \gamma = \frac{p_1 + p_2 k^2}{1 + p_3 k^2} \quad (14)$$

with $p_i = p_i(a)$

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- background level : Dark Fluid with $w(a)$
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- variable G_N , $G_N \rightarrow G_N(a, k)$
- decomposed as $G_1(a) + \text{fifth force}$

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Simple but powerful approach ! Newtonian **spherical** motion

$$\ddot{\mathbf{r}} = -\frac{\ddot{a}}{a}\mathbf{r} - \nabla\Phi \quad (15)$$

With the Poisson equation

$$k^2\Phi = -\mu(a, k)\delta\rho \quad \xRightarrow{FT} \quad \nabla^2\Phi = [\delta\rho * \tilde{\mu}] \quad (16)$$

$\tilde{\mu}$ can be written :

$$\tilde{\mu}(r) = \tilde{G}\delta_D^{(3)}(r) + (G_0 - \tilde{G})\frac{1}{4\pi R^2 r}e^{-\frac{r}{R}} \quad (17)$$

where

$$G_0 = \frac{1}{\rho_4}, \quad \tilde{G} = \frac{\rho_3}{\rho_5}, \quad R^2 = \frac{\rho_5}{\rho_4} \quad (18)$$

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$$k^2\Phi = -\mu(a, k)\delta\rho \quad \xRightarrow{FT} \quad \nabla^2\Phi = [\delta\rho * \tilde{\mu}] \quad (16)$$

$\tilde{\mu}$ can be written :

$$\tilde{\mu}(r) = \tilde{G}\delta_D^{(3)}(r) + (G_0 - \tilde{G})\frac{1}{4\pi R^2 r}e^{-\frac{r}{R}} \quad (17)$$

where

$$G_0 = \frac{1}{\rho_4}, \quad \tilde{G} = \frac{\rho_3}{\rho_5}, \quad R^2 = \frac{\rho_5}{\rho_4} \quad (18)$$

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3 radial regimes :

- 1 $r \gg R(a)$: fully GR regime (usual GR), $G_N = G_0(a)$
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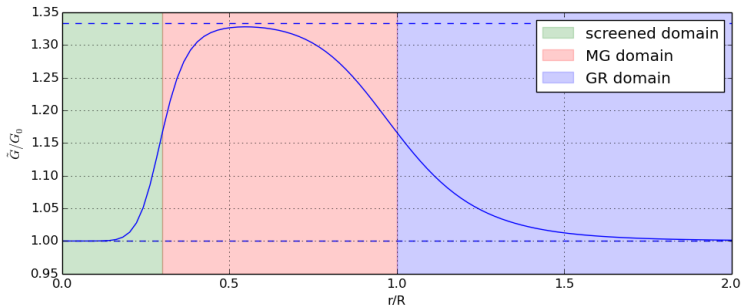
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introducing $\psi(t) = \chi(t)/\chi_0$, the *e.o.m.* for *each shell* becomes

$$\frac{\partial^2 \psi}{\partial \theta^2} + \frac{1}{\sqrt{2\Omega_m}} \frac{\partial \psi}{\partial \theta} = G_{\text{eff}} \left[\psi - \frac{1 + \Delta_0}{\psi^2} \right] \quad (19)$$

where $d\theta = \sqrt{\Omega_m/2} d \log(a)$ and $G_{\text{eff}}(r_0/R, G_0, \tilde{G})$ (in GR, $G_{\text{eff}} = 1$)

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We know the dynamics. We need initial realistic matter profiles !

Possible answer

Averaged large scale matter profiles around ... Halos and Cosmic Voids !

- analytically predicable (Gaussian primordial field)
- spherically averaged
- suitable for spherical collapse
- observable

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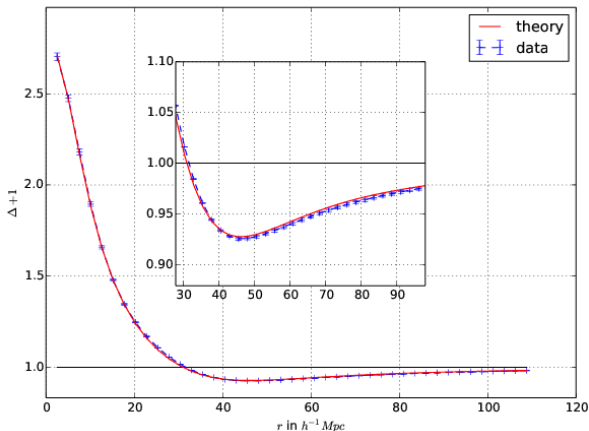


Figure : Mass contrast for $R_1 = 30 h^{-1} \text{Mpc}$

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A Modified gravity model ! Generic effective model : $f(R)$ Hu & Sawicki model (2007)

$$S = \frac{1}{2\kappa} \int d^4x \sqrt{-g} (R + f(R)) + S_m[\Psi; g] \quad (20)$$

$$f(R) = f_0 \frac{\gamma + 1}{\gamma + R_0/R} \quad (21)$$

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$$f(R) = f_0 \frac{\gamma + 1}{\gamma + R_0/R} \quad (21)$$

Spherical collapse in Modified Gravity

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initial conditions (averaged profiles) the dynamics for every model

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Hu & Sawicki model

- effective $w(a) \simeq w_0 - \frac{8}{3}|f_R^0|(a-1)$

- scalaron scale

$$R(a) \sim |f_R^0| a^{5/2} \times \left(1 + [1 - 3w(a)] \frac{1 - \Omega_m}{\Omega_m} \right)^{-3/2}$$

- usual $G_0 = 1$

- effective G in MG regime : $\tilde{G} = 4/3 G_0$

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Mass profiles for halos

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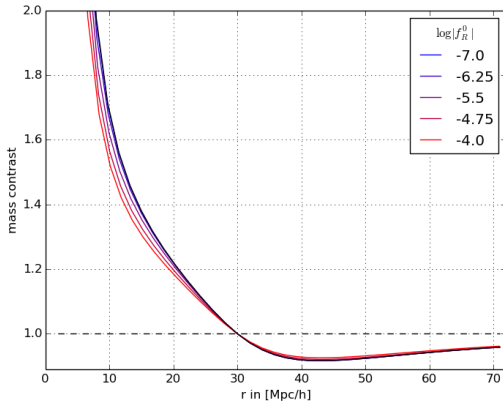


Figure : Mass contrast with various f_R^0 (Λ CDM for $f_R^0 \rightarrow 0$)

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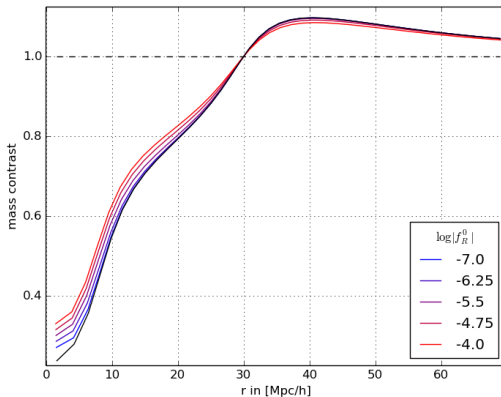


Figure : Mass contrast with various f_R^0 (Λ CDM for $f_R^0 \rightarrow 0$)

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Exclusion contour

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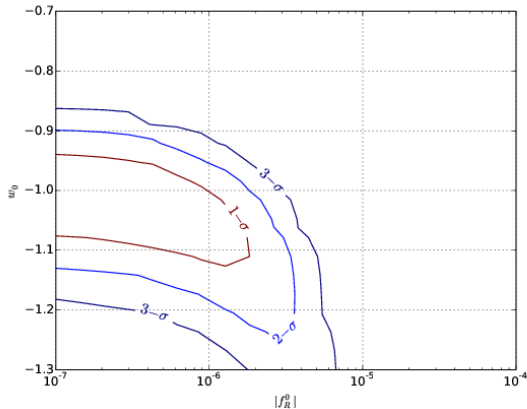


Figure : Exclusion region for w_0 and f_R^0

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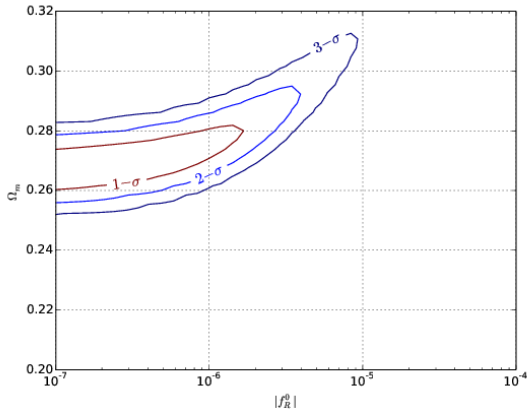


Figure : Exclusion region for Ω_m and f_R^0

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- Modified gravity and Quintessence \Rightarrow Dark Energy
- for background : $MG \equiv$ Quintessence
- Impact on LSS
- Using averaged profiles : new constraints on MG !

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Thank you for your attention !



“But before we move on, allow me to
belabor the point even further...”